Specific Heat of a Solid – Ionic Contribution Debye Model

Classical Law of Dulong & Petit

6N harmonic degrees of freedom – $\begin{cases} 3N \text{ momentum} \\ 3N \text{ normal coords} \end{cases}$

$$C_V = (6N)\left(\tfrac{1}{2}k_B\right) = 3Nk_B \quad \Rightarrow \quad \frac{C_V}{V} = 3k_B n \qquad n = \frac{N}{V}$$

In QM treatment, the 3N momenta + 3N normal coords can be thought of as 3N harmonic oscillators. These oscillations are the sound waves of vibration in the solid. We can approx their dispersion relation as

$$\omega = c_s |\vec{k}| \qquad \vec{k} \text{ is wave vector}$$

3 polarizations at each $\vec{k}$: $s = \begin{cases} L & \text{longitudinal mode, ion displacement} \parallel \vec{k} \\ T_1, T_2 & \text{transverse modes, ion displacement} \perp \vec{k} \end{cases}$

For a solid of volume V, the only sound modes are those that obey periodic boundary conditions

$$\mu = x, y, z \qquad k_\mu L = 2\pi n_\mu \qquad n_\mu = 0, 1, 2, \ldots \text{ integer}$$

ie $\frac{L}{\lambda} = n$ integer

$$\vec{k} = \frac{2\pi}{L}\vec{m}$$

The total number of sound modes = total number of oscillators = 3N. This sets an upper bound on $|\vec{k}|$

$$3N = \underbrace{3}_{\substack{\text{3 polarizations for} \\ \text{each } \vec{k}}} \sum_{|\vec{k}| < k_{max}} 1 = 3\frac{V}{(2\pi)^3} \int_{|\vec{k}| < k_{max}} d^3k = \frac{3V}{(2\pi)^3}\int_0^{k_{max}} dk\, 4\pi k^2$$

assume all 3 polarizations have same sound speed c_s

$$3N = \sum_{s} \sum_{\vec{k}} = 3 \sum_{|\vec{k}| < k_{max}} = 3 \frac{V}{(2\pi)^3} \int_{|\vec{k}| < k_{max}} d^3 k$$

$$\frac{1}{(\Delta k)^3} = \frac{V}{(2\pi)^3}$$

$$\Rightarrow 3N = \frac{3V}{(2\pi)^3} \int_0^{k_{max}} dk \, 4\pi k^2 = \frac{3V}{(2\pi)^3} \frac{4\pi}{3} k_{max}^3$$

ion density $n = \frac{N}{V} = \frac{k_{max}^3}{6\pi^2}$

Call $\omega_{max} = c_s k_{max}$ $\qquad \omega_{max}$ = Debye freq

$k_B \Theta_D = \hbar \omega_{max}$ $\qquad \Theta_D$ = Debye temperature

$$\omega_{max}^3 = 6\pi^2 c_s^3 n \qquad \Theta_D = \frac{\hbar c_s}{k_B} (6\pi^2 n)^{1/3} \sim 200\text{-}300^\circ K$$

Now $\langle E \rangle = \sum_{s} \sum_{\vec{k}} \hbar \omega_s(k) \left[\langle n_{sk} \rangle + \frac{1}{2} \right]$

boson occupation factor

$$C_V = \frac{\partial \langle E \rangle}{\partial T} = \sum_s \sum_k \frac{\partial}{\partial T} \left(\frac{\hbar \omega_s(k)}{e^{\beta \hbar \omega_s(k)} - 1} \right)$$

$$= \frac{\partial}{\partial T} \left(\frac{3 \hbar c_s}{(2\pi)^3} V \int_0^{k_{max}} dk \, 4\pi k^2 \frac{k}{e^{\beta \hbar c_s k} - 1} \right)$$

$$= \frac{\partial}{\partial T} \left(\frac{3 \hbar c_s V}{2\pi^2} \int_0^{k_{max}} dk \frac{k^3}{e^{\beta \hbar c_s k} - 1} \right)$$

$$\frac{C_V}{V} = \frac{3\hbar c_s}{2\pi^2} \int_0^{k_{max}} dk \; \frac{k^3 \left(\frac{\hbar c_s k}{k_B T^2}\right) e^{\beta\hbar c_s k}}{\left(e^{\beta\hbar c_s k}-1\right)^2}$$

$$= \frac{3\,(\beta\hbar c_s)^2 k_B}{2\pi^2} \int dk \, \frac{k^4 e^{\beta\hbar c_s k}}{\left(e^{\beta\hbar c_s k}-1\right)^2}$$

$$= \frac{3 k_B}{2\pi^2}\left(\frac{1}{\hbar c_s \beta}\right)^3 \int_0^{x_{max}} dx \, \frac{x^4 e^x}{(e^x-1)^2} \qquad x_{max} = \frac{\hbar c_s k_{max}}{k_B T} = \frac{\theta_D}{T}$$

use $\theta_D = \frac{\hbar\omega_{max}}{k_B} \Rightarrow \theta_D^3 = \frac{\hbar^3 \omega_{max}^3}{k_B^3} = \frac{\hbar^3 6\pi^2 c_s^3 n}{k_B^3}$

$$\Rightarrow \frac{\hbar^3 c_s^3}{k_B^3} = \frac{\theta_D^3}{6\pi^2 n}$$

$$\frac{C_V}{V} = \frac{3}{2\pi^2} k_B \, 6\pi^2 n \left(\frac{T}{\theta_D}\right)^3 \int_0^{\theta_D/T} dx \, \frac{x^4 e^x}{(e^x-1)^2}$$

$$\frac{C_V}{V} = 9 n k_B \left(\frac{T}{\theta_D}\right)^3 \int_0^{\theta_D/T} dx \, \frac{x^4 e^x}{(e^x-1)^2}$$

limits: $T \to \infty$ so θ_D/T is small, integrand can be expanded for small x

$$\frac{x^4 e^x}{(e^x-1)^2} \approx \frac{x^4}{x^2} = x^2$$

$$\int_0^{\theta_D/T} dx \, x^2 \simeq \frac{1}{3}\left(\frac{\theta_D}{T}\right)^3 \Rightarrow \frac{C_V}{V} = 9 n k_B \left(\frac{T}{\theta_D}\right)^3 \cdot \frac{1}{3}\left(\frac{\theta_D}{T}\right)^3$$

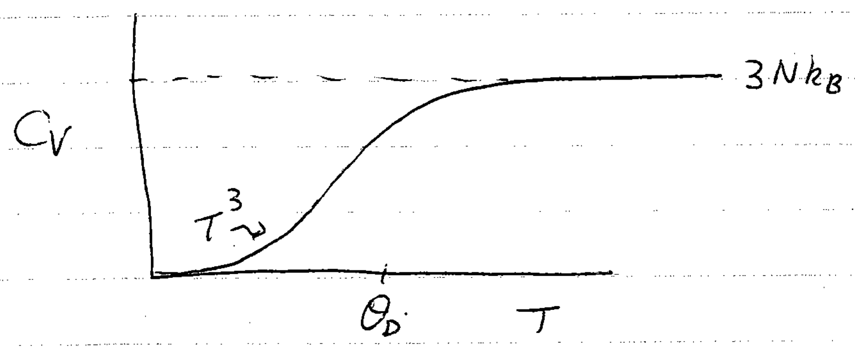
$= 3 n k_B$ Classical result

low T

$$\frac{C_V}{V} \simeq 9 n k_B \left(\frac{T}{\Theta_D}\right)^3 \underbrace{\int_0^\infty dx \frac{x^4 e^x}{(e^x - 1)^2}}_{\text{do integral}}$$

$$= \frac{12 \pi^4}{5} n k_B \left(\frac{T}{\Theta_D}\right)^3 \qquad \frac{4}{15}\pi^4$$

$\frac{C_V}{V} \propto T^3$ at low temperatures



Originally Einstein treated the problem quantum mechanically assuming constant frequencies of vibration

$\omega = \omega_0$ all modes.

Gave exponentially decreasing C_V at low T.
Debye mode is more physical.

Black Body Radiation

Cavity radiation — a volume V at fixed temp T absorbs + emits electromagnetic radiation. What are characteristics of this equilib radiation at fixed T?

EM waves with wave vector $\vec{k}$, freq $\omega = c|\vec{k}|$
two transverse polarizations for each $\vec{k}$.

Regard each mode as an oscillator. If excited to energy level n, the energy in the oscillator is

$$\epsilon = n\hbar\omega = n\hbar c k \Rightarrow n \text{ "photons" in this mode}$$

average energy in a given mode is therefore

$$\langle \epsilon \rangle = \hbar\omega \langle n \rangle = \frac{\hbar\omega}{e^{\beta\hbar\omega} - 1}$$

(ignore ground state energy $\frac{1}{2}\hbar\omega$ as it is T-indep constant)

For a volume $V = L^3$, periodic boundary conditions give the allowed wave vectors $\vec{k} = \frac{2\pi}{L}\vec{n}$ $\quad n_x, n_y, n_z$ integers

Density of states $g(\omega)$

two polarizations for each $\vec{k}$

$$\int g(\omega)\, d\omega = 2 \sum_{\vec{k}} = \frac{2V}{(2\pi)^3} \int d^3k$$

$$\Rightarrow g(\omega)\, d\omega = \frac{2V}{(2\pi)^3} 4\pi k^2 dk = \frac{V}{\pi^2} \frac{\omega^2 d\omega}{c^3}$$

$$g(\omega) = \frac{V\omega^2}{\pi^2 c^3}$$

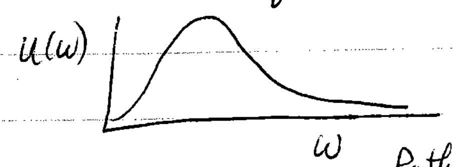
average energy per volume at freq ω is

modes at freq ω

average energy in a given mode at freq ω

$$u(\omega) = \frac{g(\omega)}{V}\left(\frac{\hbar\omega}{e^{\beta\hbar\omega}-1}\right)$$

$$u(\omega) = \frac{\hbar\,\omega^3}{\pi^2 c^3 (e^{\beta\hbar\omega}-1)}$$ ← Black Body Spectrum Plancks formula



Pathria fig 7.7

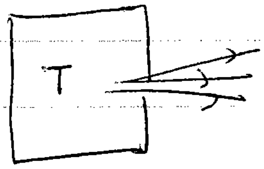
<u>Total energy density</u>

$$\frac{U}{V} = \int_0^\infty u(\omega)\,d\omega = \frac{\hbar}{\pi^2 c^3}\int_0^\infty d\omega\,\frac{\omega^3}{e^{\beta\hbar\omega}-1}$$

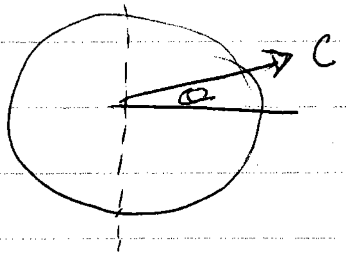
$$= \frac{\hbar}{\pi^2 c^3}\frac{1}{(\beta\hbar)^4}\underbrace{\int_0^\infty dx\,\frac{x^3}{e^x-1}}_{\frac{\pi^4}{15}} \qquad x = \beta\hbar\omega$$

$$\frac{U}{V} = \left(\frac{\pi^2 k_B^4}{15\,\hbar^3 c^3}\right)T^4$$

energy flux from a cavity, exiting from a hole



$$\text{flux } \mathcal{F} = \underbrace{\left(\frac{U}{V}\right)}_{\text{energy density}} \underbrace{c}_{\text{speed}} \underbrace{\langle\cos\theta\rangle}_{\text{projection of velocity in outwards direction}}$$



$$\langle\cos\theta\rangle = \frac{1}{4\pi}\int_0^{2\pi} d\phi \int_0^{\pi/2} d\theta \sin\theta\cos\theta \quad \leftarrow \text{only in outward direction}$$

$$= \frac{2\pi}{4\pi}\left(\frac{\sin^2\theta}{2}\right)_0^{\pi/2} = \frac{1}{4}$$

$$\mathcal{F} = \left(\frac{U}{V}\right)\frac{c}{4} = \sigma T^4 \leftarrow \text{Stefan Boltzman Law}$$

$$\text{where } \underbrace{\sigma}_{\text{Stefan's constant}} = \frac{\pi^2 k_B^4}{60\,\hbar^3 c^2} = 5.7\times 10^{-8}\ \frac{W}{m^2\cdot {}^\circ K^4}$$

We also have

$$\frac{pV}{k_B T} = \ln \mathcal{Z} = -\sum_k \underbrace{2}_{\text{polarizations}} \ln\left(1 - e^{-\beta\epsilon_k}\right)$$

$$= -\frac{2V}{(2\pi)^3}\int dk\, 4\pi k^2 \ln\left(1 - e^{-\beta\hbar c k}\right)$$

$$= -\int_0^\infty d\omega\, g(\omega) \ln\left(1 - e^{-\beta\hbar\omega}\right)$$

$$= -\frac{V}{\pi^2 c^3}\int_0^\infty d\omega\, \omega^2 \ln\left(1 - e^{-\beta\hbar\omega}\right)$$

integrate by parts

$$\frac{pV}{k_BT} = \frac{-V}{\pi^2 c^3}\Big[\underbrace{\frac{\omega^3}{3}\ln\left(1-e^{-\beta\hbar\omega}\right)\Big]_0^\infty}_{=\,0} + \frac{V}{\pi^2 c^3}\int d\omega\,\frac{\omega^3}{3}\,\frac{\beta\hbar e^{-\beta\hbar\omega}}{1-e^{-\beta\hbar\omega}}$$

$$\frac{pV}{k_BT} = \frac{V\beta\hbar}{3\pi^2c^3}\int_0^\infty d\omega\left(\frac{\omega^3}{e^{\beta\hbar\omega}-1}\right)$$

compare with computation of $\frac{U}{V}$

$$= \frac{\beta}{3}U = \frac{1}{3}\frac{U}{k_BT}$$

$$\Rightarrow \boxed{\frac{1}{3}U = PV} \quad \text{pressure of photon gas}$$

compare to non relativistic ideal gas

$U = \frac{3}{2}Nk_BT$, $PV = Nk_BT \Rightarrow \frac{2}{3}U = PV$